

Questions and Answers

The Formation of Stars and Planetary Systems, 2010, September 6-9, Särö, Sweden

Section & Talk by P. Cambré et al / A. Stutz

Name/Question N. Calvet

Do you have (near IR spectra to identify 3μ feature
can you obtain)

predicted in the OH opacity?

Name/Answer A. Stutz

We do not have such a spectrum. Need to calculate the sensitivities,
and to see if such a detection is possible.

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Section & Talk by *Lamhardt / Stutz*

Name/Question *V. Di-Hoehn*

*Do you also have an image of the N_2H^+ line in B335
to complete the textbook image ?*

Name/Answer *A. Stutz*

*No. Maps should exist, but we have not obtained
such a map ourselves, yet.*

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Section & Talk by Lamhardt / Stutz

Name/Question F. Heitsch

What is the motivation for using a Gaussian profile for the 3rd dimension to reconstruct the temperature & density structure? Would a deconvolved observed profile make a significant difference?

Name/Answer R. Lamhardt

Using a gaussian profile was simply motivated by testing the method with simplest assumptions.

Deriving deconvolved sky-plane profiles and applying those (or the closest model) to the profile in the 3rd direction is clearly the next step.

While the shape of the profile has an effect on the absolute value of the central density, ~~that~~ tests with varied widths of the gaussian profile showed that the effect on the estimate of midplane temperature is very small.

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Section & Talk by R. Launhardt / ASMK

Name/Question P. Gsell:

What is the estimated uncertainty in β ? Did you find any correlation between β and (e.g.) volume density of the objects?

Name/Answer R. Launhardt

The pure fitting uncertainty is of order ± 0.1 . However, extended background recovery introduces larger systematic uncertainties that might go up to ± 0.3 . We will have to investigate this more systematically and find an optimized and reliable approach to how filamentary background emission in, e.g., W-pass filtered PACS images, is treated. For these reasons we currently also do not want to draw any conclusions on correlations between β and, e.g., density and temperature. We are confident that we get the systematics down to $\sigma_\beta \sim 0.1$ and then we will also look for such correlations.

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Section & Talk by R. Lammhardt/A. Stutz

Name/Question N. Evans

You showed constraints on β from 100 - 1000 μm .
Models suggest that β would change over that range. Have tried breaking it up into ranges, such as 100 - 350 μm and 350 - 1000 μm ?

Name/Answer R. Lammhardt

We have not done this yet because we still try to get a better handle on calibration and by-recovering differences between different instruments. We will look into this once we have a better handle on these issues, but I am not sure if the data will allow to draw reliable conclusion on a possible β range in the FIR