

Questions and Answers

The Formation of Stars and Planetray Systems, 2010, September 6-9, Särö, Sweden

Section & Talk by: Bart Vandenbussche

Name/Question: Bruno Merin

You report the detection of [OI] and [CII] lines in Beta Pictoris. Is there any other one?

Answer: It is too early to assess that but we are working on it.

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Name/Question Alexander Krivov

Your $\beta=0.34$ result would imply grains that are larger than the wavelength, whereas the detection of the forsterite feature would imply the opposite, i.e. grains smaller than λ . What do you think the explanation could be?

Name/Answer Bart Vandenbussche

The forsterite we detect in the MIR and at $69\mu\text{m}$ is not necessarily the same small, hot crystalline silicates seen with SUBARU close to the sky. We need to look at scenarios where the small silicates are blown away on eccentric orbits and potentially funnelling, colder population.

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Name/Question Jonh Maddison

Do we know the α SED
slope of the SW blob?

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Name/Answer Bart Vandenbussche

Not yet. ~~But~~ In the spire maps
the PSF blends the blob with the
disk, we are working on deconvolving the
images to be able to determine the color
of the blob.

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Name/Question Karl Stapelfeldt

- Did you see any off-source or extended $[O\text{I}]$ emission near β Pic?
- In the 12 years since the original JCMT image, β Pic's proper motion has moved the system by $3''$. Can you tell if the submm blob to the SW of the star is comoving or not?
- How can the community get access to the improved PACS PSF data to be released shortly?

Name/Answer